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# Origin of temperature anisotropies in the cold plasma sheet: Geotail observations around the Kelvin-Helmholtz vortices

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**Abstract.** To further our understanding of the solar wind entry across the magnetopause under northward IMF, we perform a case study of a duskside Kelvin-Helmholtz (KH) vortex event on 24 March 1995. We have found that the protons consist of two separate (cold and hot) components in the magnetosphere-like region inside the KH vortical structure. The cold proton component occasionally consisted of counter-streaming beams near the current layer in the KH vortical structure. Low-energy bidirectional electron beams or flat-topped electron distribution functions in the direction along the local magnetic field were apparent on the magnetosphere side of the current layer. We discuss that the bidirectionality of the electrons and the cold proton component implies magnetic reconnection inside the KH vortical structure. In addition, we suggest selective heating of electrons inside the vortical structure via wave-particle interactions. Comparing temperatures in the magnetosphere-like region inside the vortical structure with those in the cold plasma sheet, we show that further heating of both the electrons and the cold proton component is taking place in the cold plasma sheet or on the way from the vortices to the cold plasma sheet.

**Keywords.** Magnetospheric physics (Magnetotail; Magnetotail boundary layers; Plasma sheet)

## 1 Introduction

The plasma sheet in the near-Earth magnetosphere becomes cold under the northward interplanetary magnetic field (IMF) (e.g. Zwolakowska et al., 1992; Zwolakowska and Popielawska, 1992; Terasawa et al., 1997; Borovsky et al., 1998; Nishino et al., 2002; Wing et al., 2005). Such cold plasma is thought to be of solar wind origin and to come into the near-Earth plasma sheet through the low-latitude

boundary layer (LLBL) in the tail-flank regions (e.g. Terasawa et al., 1997; Fujimoto et al., 1998). The entry mechanism of the solar wind plasma under northward IMF is still under debate.

As candidates for entry mechanism of the solar wind plasma under northward IMF, double high-latitude reconnection on the dayside and Kelvin-Helmholtz (KH) instability in the tail flanks have been proposed. Double high-latitude reconnection has been thought to take place under strongly northward IMF (Song and Russell, 1992; Li et al., 2005; Øieroset et al., 2005), and in-situ evidence of this process was found by recent satellite observations (Onsager et al., 2001; Lavraud et al., 2005, 2006). The KH instability is expected to occur in the flank regions because of velocity shear across the magnetopause (Fairfield et al., 2000). Simulations (Otto and Fairfield, 2000; Nykyri and Otto, 2001; Matsumoto and Hoshino, 2006; Nakamura et al., 2006), as well as observational studies (e.g. Hasegawa et al., 2004, 2006) have suggested a development of the vortical structure under strongly northward IMF. Furthermore, simulation studies (e.g. Otto and Fairfield, 2000; Nykyri and Otto, 2001; Nakamura et al., 2006) suggested that magnetic reconnection may occur in the KH vortices and that it can play an important role in plasma transport across the magnetopause. A recent observational study by Nykyri et al. (2006) suggested reconnection signatures in the flank region on the dawnside associated with KH instability, although the IMF did not point strongly northward in their event.

In the cold plasma sheet and the LLBL, the electrons frequently have a parallel anisotropy where parallel temperature dominates over the perpendicular one (Hada et al., 1981; Traver et al., 1991; Phan et al., 1997; Fujimoto et al., 1998). The parallel anisotropy of low energy protons in the dusk flank plasma sheet has also been pointed out by Traver et al. (1991) and Nishino et al. (2007a). A recent study by Nishino et al. (2007b) showed that strong parallel anisotropies of both the electrons and the cold proton component occur under

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strongly northward IMF, and suggested that the anisotropies may be related to the rolled-up KH vortices in the tail flank. Nishino et al. (2007b) also found that the low-energy portion of the electron distribution functions in the cold plasma sheet has a flat-topped shape, and that the parallel anisotropy of electrons is stronger than that of the cold proton component, which implies that some other mechanism, such as wave-particle interactions, works besides adiabatic heating.

In order to study strong parallel anisotropies around the KH vortices under strongly northward IMF, we focus on the duskside magnetopause crossing event on 24 March 1995. During this day strongly northward IMF was monitored by upstream Wind for a prolonged interval, and the Geotail spacecraft observed quasi-periodic wavy structures around the LLBL on the duskside (Fujimoto et al., 1998). Fairfield et al. (2000) suggested that the wavy structures are due to the development of KH vortices, with the support of simulations by Otto and Fairfield (2000) that proposed the occurrence of magnetic reconnection in the KH vortical structures. Furthermore, the magnetopause has been shown to be rolled-up at times due to the KH instability, as demonstrated in an observational study by Hasegawa et al. (2006). In the cold plasma sheet adjacent to the LLBL, the ions consisted of two separate components, which showed spatial admixture of the solar wind ions and the magnetospheric hot ions (Fujimoto et al., 1998). The low-energy portion of the electrons on the magnetosphere side around the boundary had a parallel (bidirectional) anisotropy along the magnetic field (Fujimoto et al., 1998; Fairfield et al., 2000). Nishino et al. (2007b) reported that both the electrons and the cold proton component in the cold plasma sheet had strong parallel anisotropies in this event. A detailed study of distribution functions, in addition to temperature anisotropies around the KH vortical structures, will provide more information about the physical mechanism which has impacted the plasmas during the cold plasma sheet formation. We examine temperature anisotropies of electrons and protons inside the KH vortical structures, as well as inspect the velocity distribution functions in detail, comparing them with those in the cold plasma sheet.

## 2 Instrumentation

We use data from the three-dimensional (3-D) electron and ion distribution functions obtained every 12 s by the low energy particle (Geotail/LEP) experiment (Mukai et al., 1994). The magnetic field data are obtained every 1/16 s by the flux-gate magnetometer (Geotail/MGF) (Kokubun et al., 1994). The electrons are detected by the electron analyzer of LEP-EAe, and we use data of the low-energy mode in the range between 8.3 eV and 7.6 keV. The ion energy-per-charge analyzer of LEP-EAi detects ions between 32 eV/q and 39 keV/q, which covers most of the typical energy range of protons in the boundary layer and the cold plasma sheet.

In the present study, all of the detected positively-charged ions are assumed to be protons. Solar wind parameters obtained by the Wind spacecraft were provided via CDAWeb. The Wind data are from SWE (Ogilvie et al., 1995) and MFI (Lepping et al., 1995). We use the GSM coordinate system throughout the paper.

## 3 Calculation of moment parameters

For electrons, we obtain effective temperatures in the perpendicular and parallel directions to the local magnetic field, which are denoted as  $T_{e\perp}$  and  $T_{e\parallel}$ , respectively. Although the typical kinetic energy of photoelectrons is less than the electric potential of the spacecraft, which is lower than 10 eV in the cold plasma sheet (Ishisaka et al., 2001), some photoelectrons have energy higher than the spacecraft potential (Ueno et al., 2001b). We therefore checked the observed electron distribution functions, and utilize PSDs with energy higher than 12.86 eV to exclude photoelectrons.

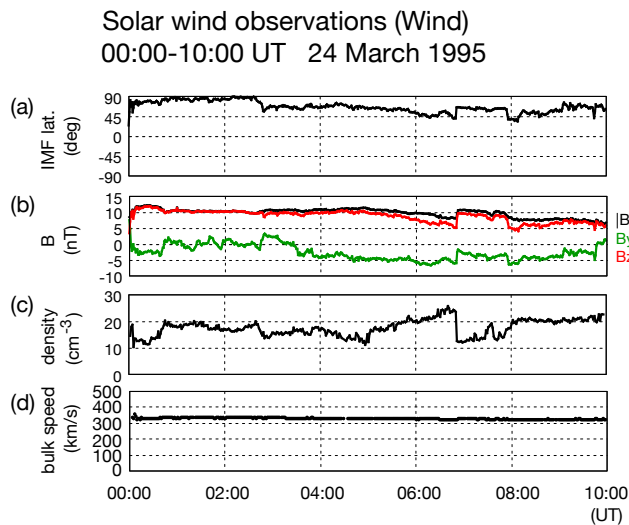
For protons, we not only calculate moment parameters by simple moment calculations but also deal with cold and hot components separately, for the data where two components are apparent. To separate the ion distribution function into the cold and hot components, we utilize a two-Maxwellian mixture model with a scheme developed by Ueno et al. (2001a), as Nishino et al. (2007a) performed. Perpendicular and parallel temperatures of the cold and hot components are denoted as  $T_{C\perp}$ ,  $T_{C\parallel}$ ,  $T_{H\perp}$ , and  $T_{H\parallel}$ , respectively. Proton temperatures from the simple moment calculations are represented by  $T_{p\perp}$  and  $T_{p\parallel}$ .

## 4 Observations

### 4.1 Overview

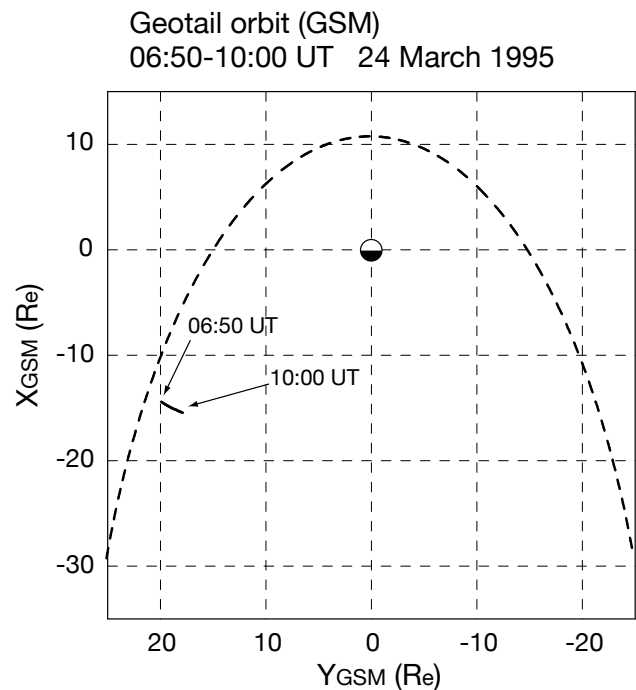
Figure 1 shows the solar wind observations between 00:00–10:00 UT on 24 March 1995. From the top, (a) latitudinal angle of the IMF ( $\theta_{\text{IMF}}$ ), (b) strength (black line),  $Y$  (green) and  $Z$  (red) components of the IMF, (c) proton density, and (d) bulk flow speed are presented. The IMF pointed strongly northward ( $\theta_{\text{IMF}} > 45^\circ$ ) and the density was higher than  $10 \text{ cm}^{-3}$  throughout the interval. The bulk flow speed was as slow as 330–340 km/s. The convection time of the solar wind from the Wind location ( $X \sim 219 R_E$ ) to the Earth's magnetosphere, which is not included in the figure, was about 70 min.

Figure 2 shows the projection onto the GSM-XY plane of the Geotail orbit between 06:50–10:00 UT. The spacecraft gradually moved toward the midnight region with its  $Z$  coordinate decreasing from  $3.9 R_E$  to  $1.9 R_E$  (not shown). Figure 3 shows an overview of the Geotail observations for the interval. From the top, (a–b) energy-time (E-t) spectrograms of omni-directional electrons and ions, (c) electron temperatures in the direction perpendicular (green) and



**Fig. 1.** Solar wind observations by Wind ( $X \sim 219 R_E$ ) between 00:00–10:00 UT on 24 March 1995. From the top, (a) the latitudinal angle of the IMF, (b) strength (black), Y (green), and Z (red) components of the IMF, (c) proton density, and (d) bulk flow speed are presented. The convection time ( $\sim 70$  min) is not included in the figure.

parallel (blue) to the local magnetic field, (d) proton temperatures, (e)  $V_X$  (blue) and  $V_Y$  (green) of bulk proton velocity, and (f) three components of the magnetic field are presented. Locations of the Geotail spacecraft are given at the bottom of the figure. Large perturbations in the magnetic field, the velocity, and the electron and proton temperatures are attributed to encounters with the vortical structures of the KH instability around the boundary. Concerning the inbound transition of the Geotail spacecraft from the magnetosheath to the plasma sheet, Fairfield et al. (2000) categorized the observed data into four separate regions: the magnetosheath (until 03:54 UT, not shown here), the exterior interaction region (between 03:54–07:16 UT), the interior interaction region (between 07:16–09:10 UT), and the cold plasma sheet (after 09:11 UT). In the interior interaction region Geotail observed magnetosheath-like dense plasma and magnetosphere-like plasma alternately with a quasi-period of about 3–5 min. By a visual inspection of the data shown in Fig. 3, we have found that around 08:00 UT the bulk flow speed slowed down, fluctuations in the magnetic field reduced, and magnetosphere-like plasma was observed for several minutes. In order to examine the transition of plasma from the dense magnetosheath-like region to the stagnant magnetosphere-like region, we study the data around 08:00 UT in more detail.



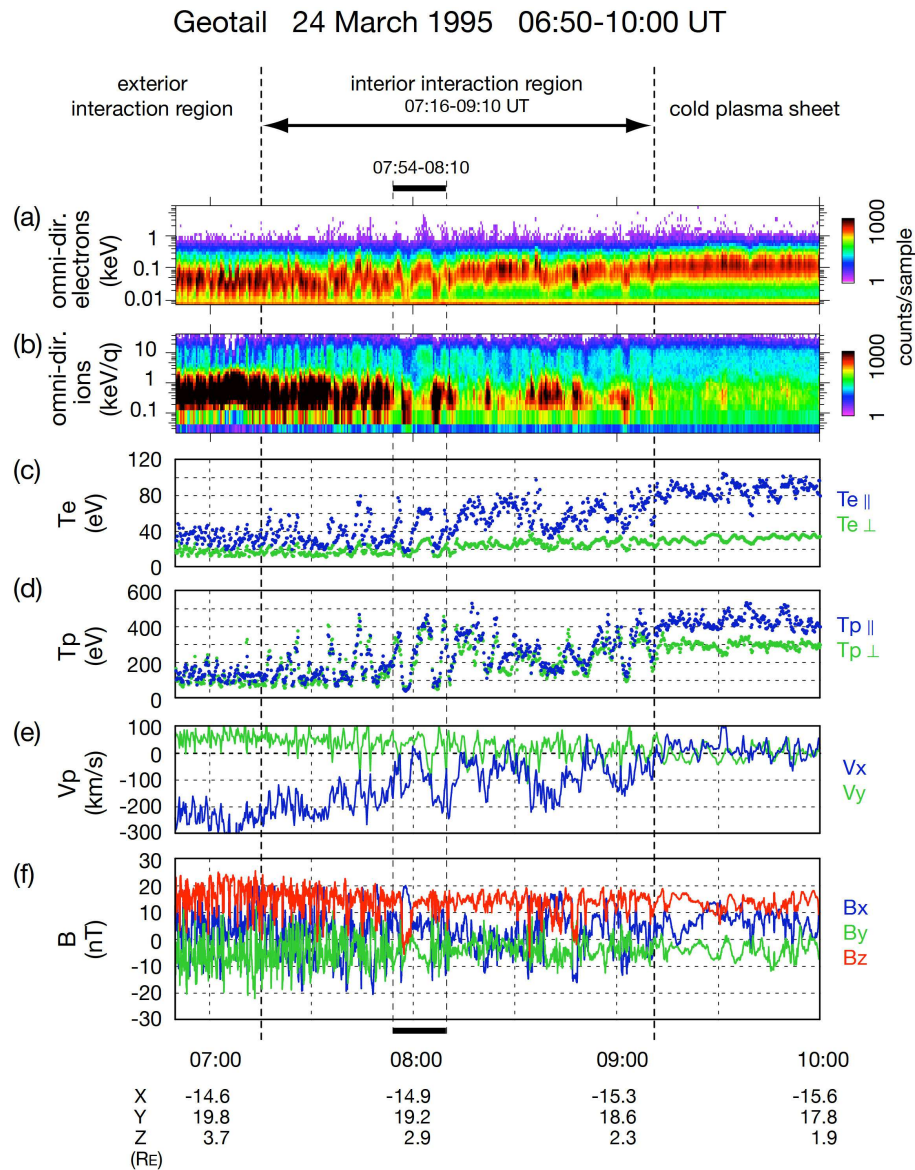
**Fig. 2.** Projection onto the GSM-XY plane of the Geotail orbit between 06:50–10:00 UT on 24 March 1995. A dotted curve in the figure corresponds to the averaged location of the magnetopause calculated with a model by Roelof and Sibeck (1993).

#### 4.2 Categorization of regions inside vortices

In Fig. 4 we expand the Geotail observation for the 16-min interval 07:54–08:10 UT during which the satellite repeated multiple crossings of the boundary. From the top, E-t spectrogram of omni-directional electrons, electron temperatures, E-t spectrograms of sunward and tailward ions, ion densities, proton temperatures from the simple moment calculations, temperatures of the hot and cold proton components, bulk proton velocity, and magnetic field are presented. At 08:00 UT Geotail was located at  $(-14.9, 19.2, 2.9) R_E$ . Prior to examining velocity distribution functions in detail, we inspect plasma moment parameters and the magnetic field, dividing the data between 07:57–08:05 UT into 6 categories, as follows.

##### (A) Magnetosheath-like region

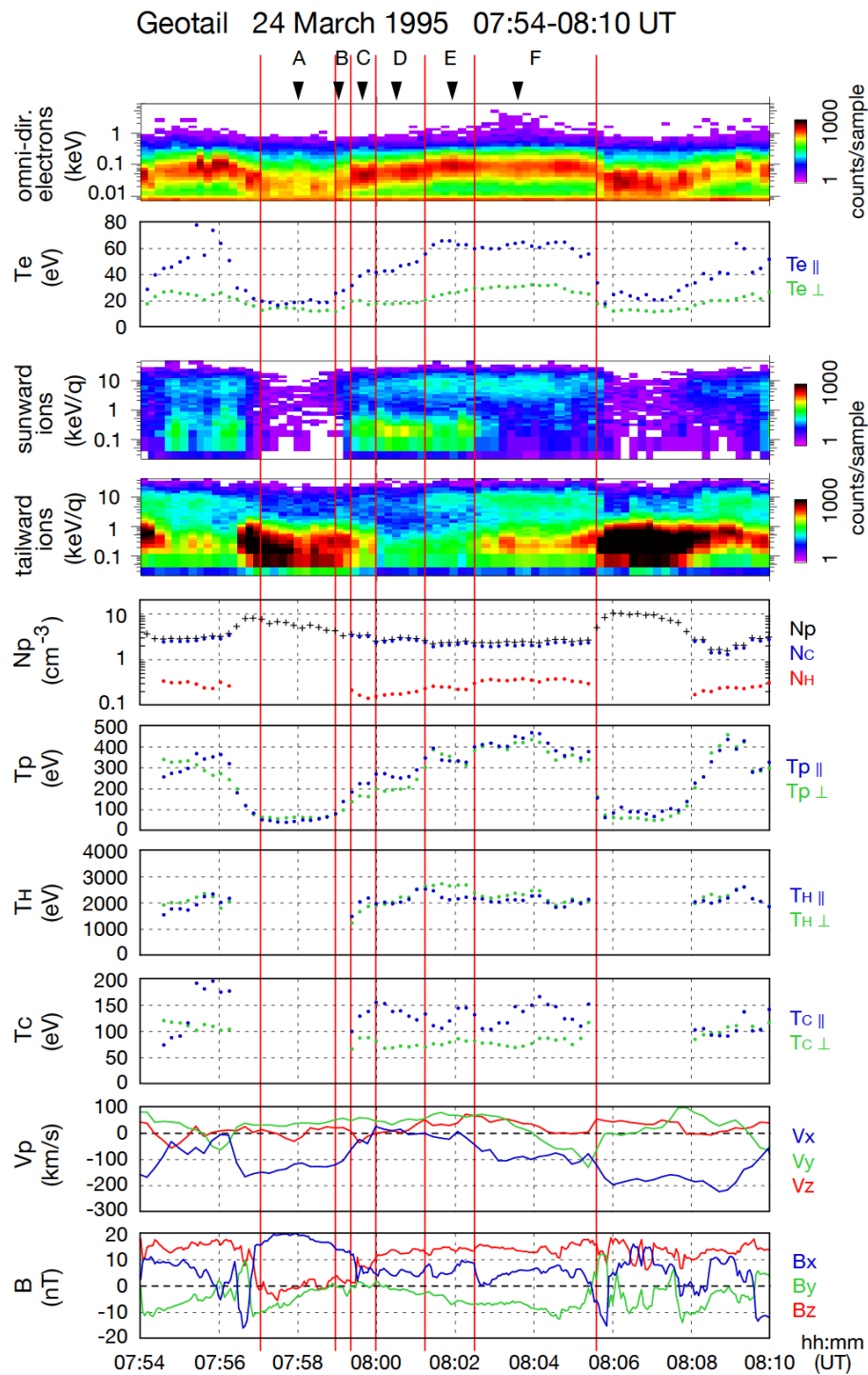
Between 07:57:05–07:58:57 UT Geotail stayed in the magnetosheath-like region, where dense ( $> 4 \text{ cm}^{-3}$ ), low-temperature plasma flows tailward ( $-150 \sim -100 \text{ km/s}$ ). Protons have a perpendicular anisotropy ( $T_{p\perp}/T_{p\parallel} \sim 1.0\text{--}1.5$ ), where  $T_{p\perp}$  and  $T_{p\parallel}$  were  $\sim 65 \text{ eV}$  and  $\sim 52 \text{ eV}$ , respectively. Since these temperatures include the effects from the higher energy component which may be from the magnetosphere, we recalculate moment parameters using ion PSD with



**Fig. 3.** An overview of the Geotail observations between 06:50–10:00 UT on 24 March 1995. From the top, (a) energy-time (E-t) spectrogram of omni-directional electrons, (b) E-t spectrogram of omni-directional ions, (c) electron temperatures in the direction perpendicular (green) and parallel (blue) to the local magnetic field, (d) perpendicular and parallel proton temperatures, (e)  $V_x$  (blue) and  $V_y$  (green) of protons in km/s, and (f) three components of the magnetic field in nT are shown. Vertical dashed lines around 07:16 UT and 09:11 UT correspond to transitions of the observed regions delineated by Fairfield et al. (2000), who categorized the data into 4 separate regions: the magnetosheath (not shown here), the exterior interaction region, the interior interaction region, and the cold plasma sheet. The observations between 07:54–08:10 UT shown by two vertical dashed lines are expanded in Fig. 4.

energy-per-charge less than 3 keV/q. (This criterion is based on a visual survey of the observed PSD.) Then we obtain a perpendicular anisotropy ( $T_{p\perp}/T_{p\parallel} \sim 1.1$ –1.7), where  $T_{p\perp} \sim 55$ –60 eV and  $T_{p\parallel} \sim 35$ –45 eV, respectively. The electron temperatures were about 15–20 eV, with the parallel temperature being slightly higher than the perpendicular temperature. In spite of the strongly northward IMF (Fig. 1a),

the magnetic field in the dense region was directed sunward, with  $B_z$  being slightly negative, which implies a complicated interaction between the magnetosheath and the magnetospheric plasma (Fairfield et al., 2000).



**Fig. 4.** Expansion of the Geotail observation between 07:54–08:10 UT on 24 March 1995. From the top, E-t spectrogram for electrons, parallel (blue) and perpendicular (green) electron temperatures, E-t spectrograms for sunward and tailward ions, proton density, parallel and perpendicular proton temperatures, temperatures of the hot and cold proton components (for selected intervals), three components of the bulk proton flow, and the magnetic field are shown. Time resolution of plasma data is 12 s and that of magnetic field data is 3 s. Black triangles at the top of the figure correspond to the data whose cuts of the velocity distribution functions are shown in Figs. 5, 7, and 8.

## (B) Boundary I

Between 07:58:57–07:59:22 UT the proton temperatures started to increase, and a large enhancement of the parallel temperature resulted in the parallel anisotropy. The parallel temperature of electrons ( $T_{e\parallel}$ ) was about 30 eV while the perpendicular temperature ( $T_{e\perp}$ ) was still  $\sim 15$  eV, which suggests parallel heating of the electrons coming from the magnetosheath-like region. The tailward bulk plasma flow decreased ( $-100$  km/s) and density decreased from  $\sim 4.3$  cm $^{-3}$  to  $\sim 3.6$  cm $^{-3}$ . The magnetic field in this region was still directed sunward, as it was in the magnetosheath-like region.

## (C) Boundary II (including a current layer)

In the boundary II between 07:59:22–07:59:59 UT the electron temperatures were further enhanced, where  $T_{e\parallel}$  increased to be 30–40 eV and  $T_{e\perp}$  was about 20 eV. Inspecting the magnetic field data at 16-Hz resolution, we found that the Geotail spacecraft crossed a current layer between 07:59:29–07:59:32 UT, which shows a sudden change in the magnetic field direction from sunward to northward. On the magnetosphere-like side of the current layer (after 07:59:34 UT), two separate (cold and hot) proton components are seen in the E-t spectrograms. The cold proton component had a parallel anisotropy ( $T_{C\perp}/T_{C\parallel} \sim 0.64$ ), where  $T_{C\perp} \sim 90$  eV and  $T_{C\parallel} \sim 130$  eV. The bulk flow was directed slightly tailward ( $-50$  km/s), and the density was  $3.5$ – $3.8$  cm $^{-3}$ . In this region the positive  $B_Z$  was not dominant but comparable to  $B_X$ .

## (D–F) Magnetosphere-like regions

After moving across the boundary layer, Geotail stayed in the magnetosphere-like region with two-component protons, where  $B_Z$  was positive and dominant. Since the two-component protons in the stagnant plasma sheet on the dusk-side is occasionally observed under northward IMF (e.g. Hasegawa et al., 2003; Wing et al., 2005; Nishino et al., 2007a), regions (D–F) resemble the cold plasma sheet (discussed later). For the two-component protons, we obtain parameters of the cold and hot components separately (Fig. 4). We divide the region into 3 categories based on observed plasma signatures, as follows.

## (D) Magnetosphere-like region I

In the initial 75 s of the plasma sheet observation (07:59:59–08:01:14 UT), Geotail was in a stagnant region. The cold proton component had a strong parallel anisotropy ( $T_{C\perp}/T_{C\parallel} \sim 0.53$ ), where  $T_{C\perp} \sim 73$  eV and  $T_{C\parallel} \sim 140$  eV. The hot proton component was almost isotropic or had a perpendicular anisotropy ( $T_{H\perp}/T_{H\parallel} \sim 0.96$ – $1.1$ ,  $T_{H\perp}$  and  $T_{H\parallel} \sim 1.9$ – $2.5$  keV). The densities of the cold and hot proton components ( $N_C$  and  $N_H$ ) were  $2.7$  cm $^{-3}$  and  $0.18$  cm $^{-3}$ , respec-

tively. The perpendicular and parallel electron temperatures ( $T_{e\perp}$  and  $T_{e\parallel}$ ) were 24 eV and 65 eV, respectively. The magnetic field was directed northward ( $B_Z \sim 13$ – $15$  nT), as is often seen in the central region of the plasma sheet.

## (E) Magnetosphere-like region II

Between 08:01:14–08:02:29 UT the parallel electron temperature was further enhanced to 60–70 eV, and the perpendicular temperature also slightly increased ( $\sim 20$ – $30$  eV). The density of the cold proton component decreased ( $N_C \sim 2.2$  cm $^{-3}$ ) while the hot proton component increased ( $N_H \sim 0.24$  cm $^{-3}$ ), where the total proton density slightly decreased but was higher than  $2$  cm $^{-3}$ . The temperatures of the cold proton component were similar to those in region (D), while the hot component had a perpendicular anisotropy ( $T_{H\perp}/T_{H\parallel} \sim 1.2$ ), where  $T_{H\perp} \sim 2.7$  keV and  $T_{H\parallel} \sim 2.2$  keV. We note that the increase in the proton temperatures ( $T_{p\perp}$  and  $T_{p\parallel}$ ) was not due to a temperature increase in the cold proton component but due to an enhancement in the fluxes of the hot proton component. In short, the electron temperatures were enhanced as the spacecraft moved from region (D) to region (E), while the temperatures of the cold proton component were not. This observation suggests that selective heating of electrons was taking place around the boundary between regions (D) and (E). In region (E) the plasma flows were low, but the positive  $V_Y$  ( $\sim 70$  km/s) was dominant.

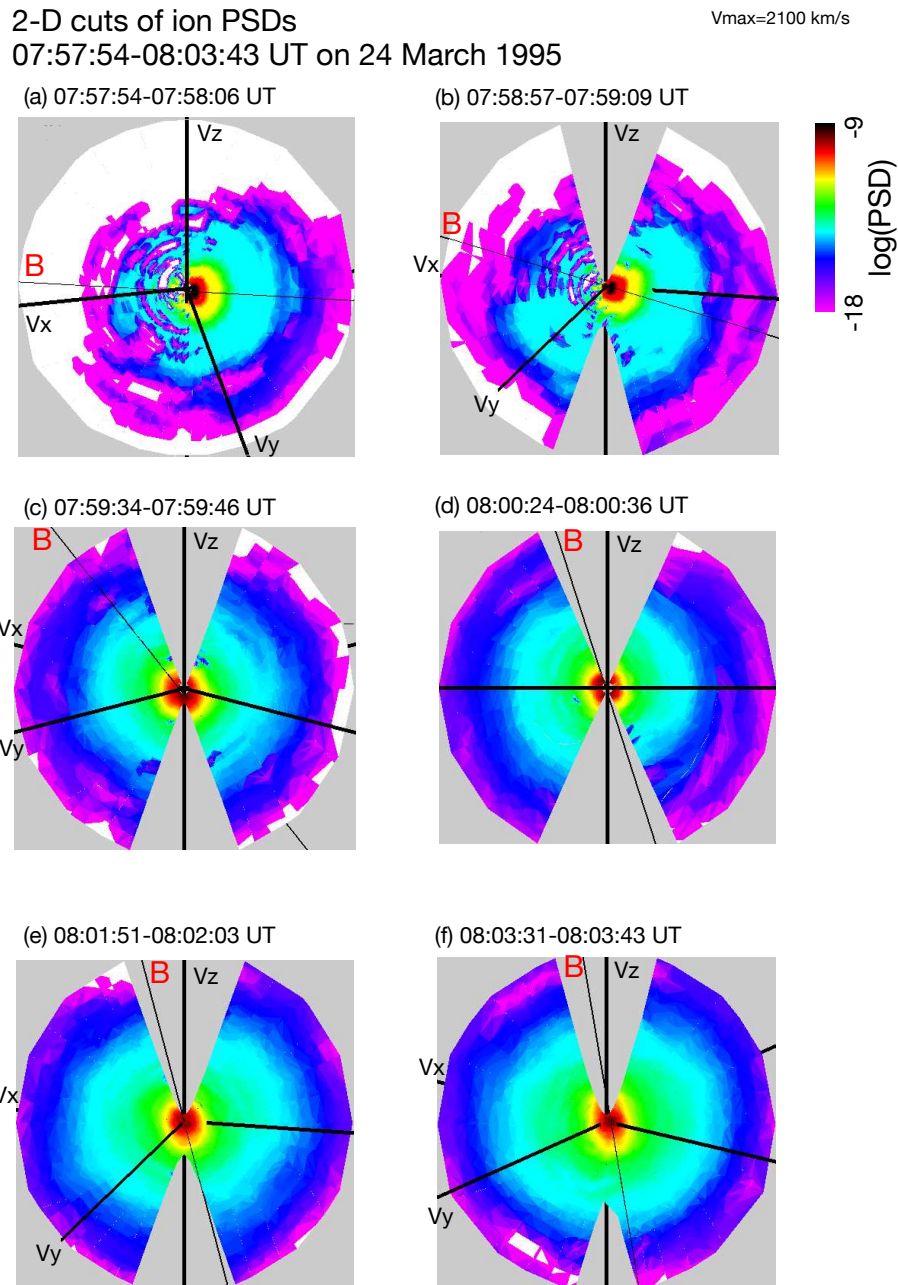
## (F) Magnetosphere-like region III

For the last 3 min of the magnetosphere-like plasma observation (between 08:02:29–08:05:36 UT), plasma flows were directed tailward ( $-100$  km/s). As for the densities,  $N_H$  increased to  $\sim 0.35$  cm $^{-3}$ , and  $N_C$  was  $\sim 2.2$  cm $^{-3}$ . Both the electrons and the cold proton component possessed strong parallel anisotropies, with temperatures being similar to those in region (E). The  $T_{e\perp}$  and  $T_{e\parallel}$  were  $\sim 30$  eV and  $\sim 61$  eV, respectively. The high-energy electrons (up to  $\sim 5$  keV), whose flux was low but undoubtedly higher than that in the ambient regions, such as (D) or (E), may be of magnetospheric origin. Geotail came out of the magnetosphere-like region around 08:05:36 UT and moved back into a dense magnetosheath-like region that resembles region (A).

## 4.3 Proton distribution functions

We examine the velocity distribution functions in each region to know the substance of the temperature anisotropies. Figure 5 presents snapshots of the ion distribution functions in each region of the 6 categories mentioned above. Each panel (a–f) shows 2-D cuts of the ion distribution function in the plane that includes the local magnetic field and the velocity obtained from the moment calculations. In region (A), the contours of the low-energy portion (red-yellow colored





**Fig. 5.** 2-D cuts of the ion distribution functions for the boundary crossing. Each panel (a–f) corresponds to categorization shown in Fig. 4. Each plane includes the local magnetic field (shown by a narrow black line and indicated by a red letter “B”) and the velocity center from the moment calculation.

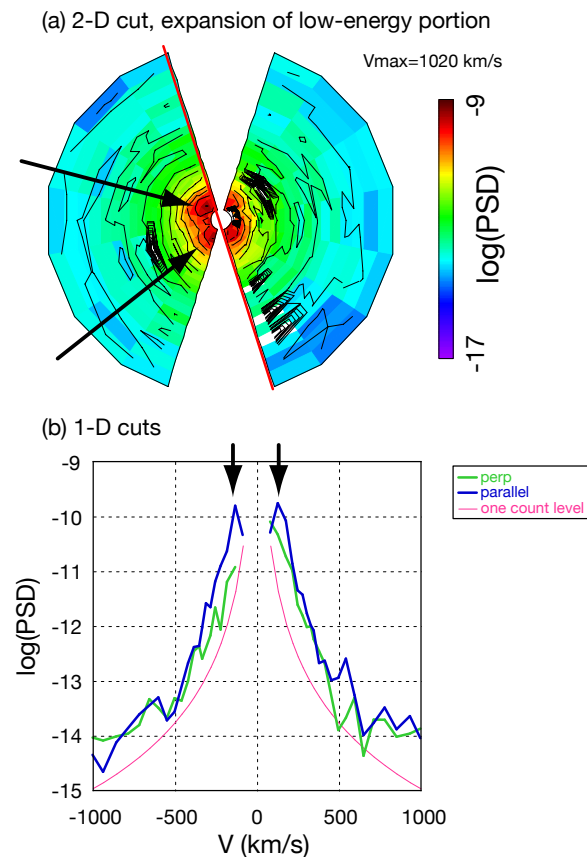
region) are elongated in the direction perpendicular to the local magnetic field. The velocity center is shifted to the right side, which corresponds to the fast tailward flow in this region. In the boundary region (B) that is on the magnetosheath side of the current layer, the elongation of the low-energy portion is weaker and the flux of the high-energy component increased. In region (C), slight elongation of the contours of the low-energy portion in the parallel direction is seen. In the

magnetosphere-like regions (D–F), further elongations of the contours in the parallel direction are apparent at low energy.

We focus on regions (D–F), where a strong parallel anisotropy of the cold proton component was observed. The top panel (a) of Fig. 6 expands the low-energy portion ( $|V| < 1020$  km/s) of the distribution function between 08:00:24–08:00:36 UT in region (D). The 1-D cuts of the distribution function in the directions perpendicular and parallel



## Counter-streaming proton beams in region (D)



**Fig. 6.** An expansion of the low-energy portion of the ion distribution function between 08:00:24–08:00:36 UT (the same data shown in Fig. 5d). The upper panel (a) shows a 2-D cut in the plane that includes the local magnetic field. Counter-streaming beams (highlighted by black arrows) along the local magnetic field are apparent. The lower panel (b) presents 1-D cuts in the directions perpendicular (green) and quasi-parallel (blue) to the magnetic field. Two red curves correspond to the one count level of the detector. Since the direction purely parallel to the magnetic field was out of sight of the detector, the phase space density in the quasi-parallel direction is presented.

to the magnetic field are shown in the bottom panel (b). Focusing on the cold proton component (red-yellow colored region), we find counter-streaming proton beams in the direction along the local magnetic field. Although the counter-streaming proton beams, which are highlighted by black arrows in the figure, are not isolated from the thermal components, the distribution function along the magnetic field has a couple of peaks in the energy range from 42 eV to 123 eV (90–150 km/s). To discuss later whether the observed proton beams fit with reconnection signature, we estimate the parallel/anti-parallel component of the magnetic field and Alfvén speed. With the parallel/anti-parallel component of

the magnetic field  $\sim 10$  nT (assumed to be  $\sim 1/\sqrt{2}$  of the magnetic field strength  $\sim 14$  nT) and the proton density  $\sim 3\text{--}4\text{ cm}^{-3}$ , we obtain the Alfvén speed ( $V_A$ ) of about 110–120 km/s. Therefore, the observed speeds of the counter-streaming proton beams are comparable to the Alfvén speed and thus consistent with the signature of magnetic reconnection. If we assume purely an anti-parallel configuration of the magnetic field on both sides of the current layer and use 14 nT as the anti-parallel field component, we obtain the Alfvén speed of  $\sim 150\text{--}180$  km/s, which is slightly faster than the observed proton beam speed but not so different from it. Such low-energy counter-streaming protons were continuously observed throughout region (D) (6 data points between 07:59:59–08:01:14 UT), while they were not seen in regions (E) nor (F) (Figs. 5e and f). In regions (E) and (F), the low-energy portion of protons consisted of a single component with a strong parallel anisotropy. In previous studies counter-streaming protons were observed near the high-latitude magnetopause (Lavraud et al., 2002; Retinò et al., 2005) and those in the dawnside low-latitude boundary (Nykyri et al., 2006). These counter-streaming protons have been interpreted as evidence of magnetic reconnection under northward IMF, and the shape of their PSDs resembles that in region (D).

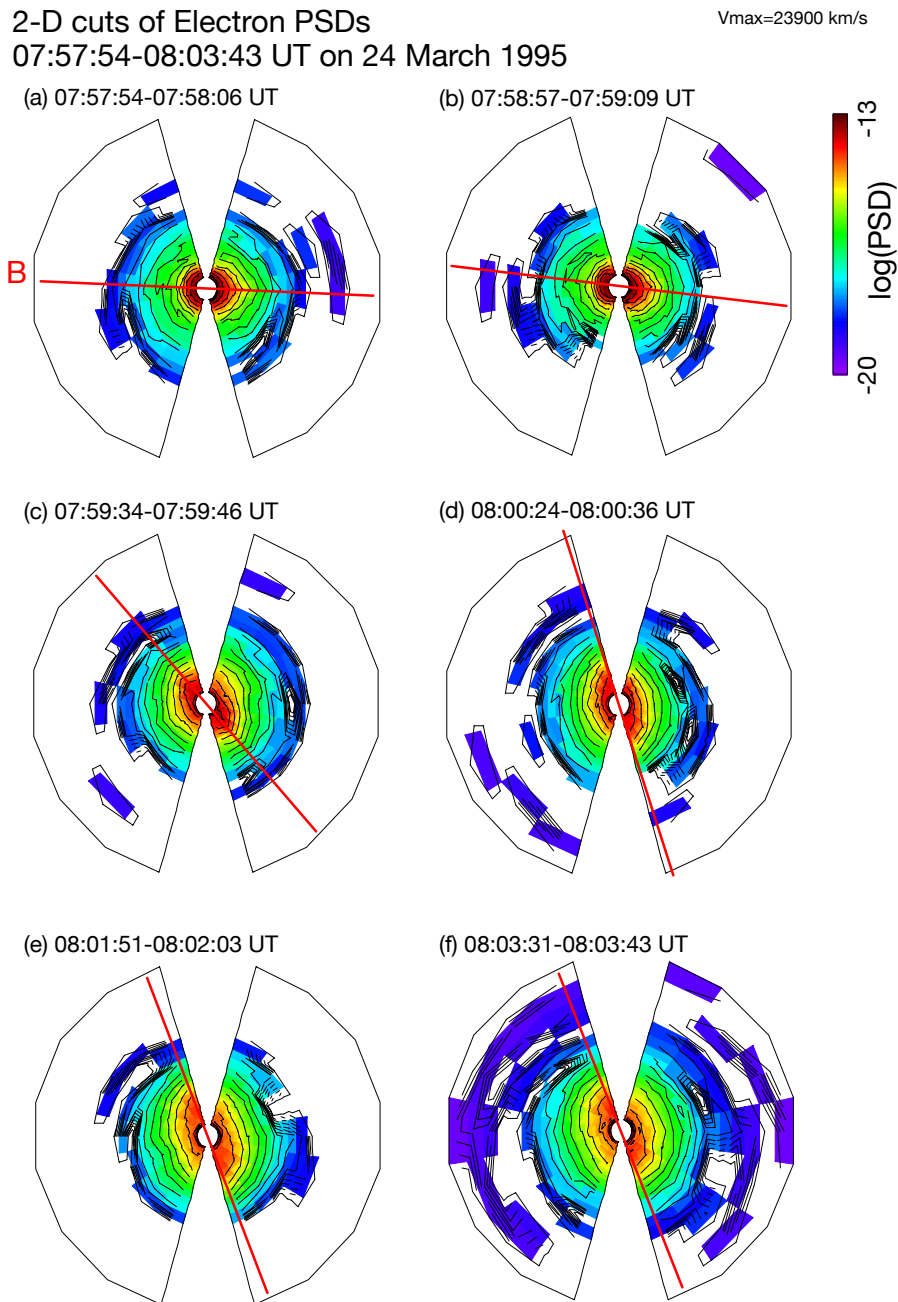
#### 4.4 Electron distribution functions

We next examine electron distribution functions in each region. Figures 7 and 8 show snapshots of 2-D and 1-D cuts of the electron velocity distribution functions for the same intervals as shown in Fig. 5. The plane shown in Fig. 7 includes the local magnetic field, and the maximum speed corresponds to about 23 900 km/s.

In the magnetosheath-like region (A), where the electrons had a weak parallel anisotropy, an enhancement of the low-energy electron flux (with speed less than 10 000 km/s) was seen only in one direction along the magnetic field and not in its opposite direction (Fig. 8a). The unidirectional enhancement of the electron flux might have originated from the electron strahl that is often observed in the solar wind flow (e.g. Rosenbauer et al., 1977; Pilipp et al., 1987), while it might also result from the occurrence of magnetic reconnection on the field line (e.g. Onsager et al., 2001).

In region (B), the parallel temperature of electrons started to be enhanced. In addition to an enhancement of the unidirectional flux in the low-energy range which may be the electron strahl of solar wind origin, the flux in the direction opposite to the strahl component was also enhanced.

Since the rotation of the magnetic field around the current layer was observed within 3 s (07:59:29–07:59:32 UT), i.e. shorter than the time resolution of the particle data ( $\sim 12$  s), we select the data right after the crossing of the current layer in the vortical structure. Figures 7c and 8c show cuts of the electron distribution function observed between 07:59:34–07:59:46 UT when the spacecraft was on the magnetosphere

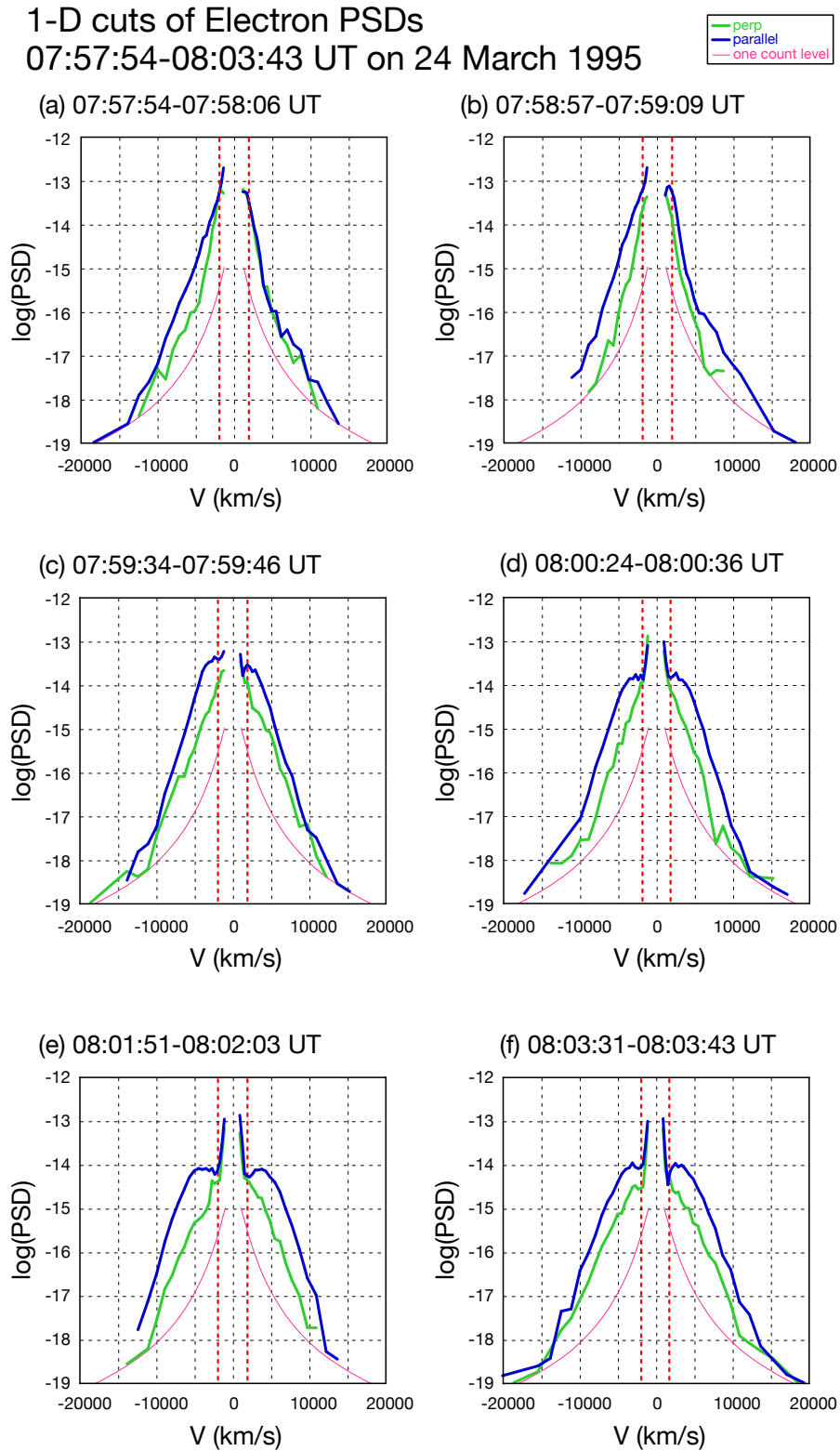


**Fig. 7.** 2-D cuts of the electron velocity distribution functions, which correspond to the ion distribution functions shown in Fig. 5.

side of the current layer. A flat-topped shape can be seen in the low-energy range (with speed less than  $\sim 4000$  km/s). The low-energy bidirectional electron component started to be observed just on the magnetosphere side of the current layer.

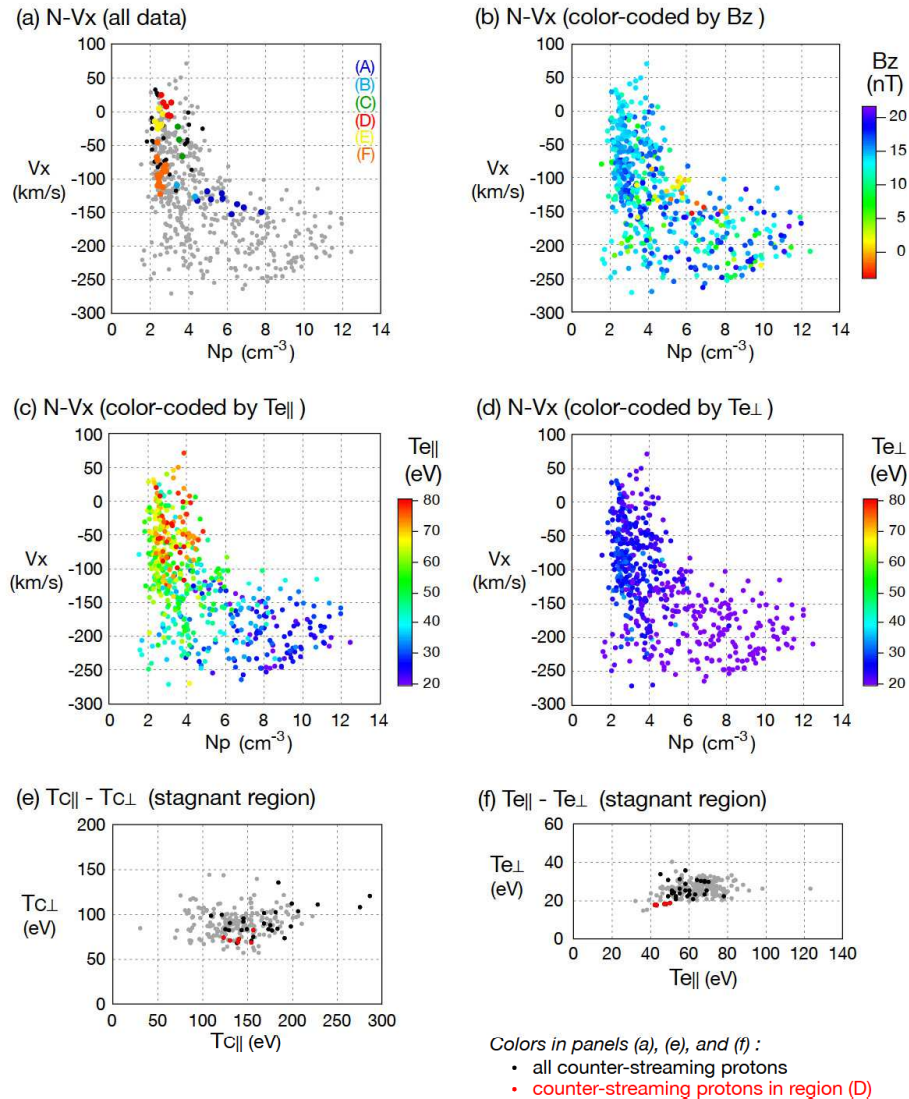
In the magnetosphere-like regions (D–F), the low-energy bidirectional component became further apparent. In region (D), we find that the electron distribution function has a flat-topped shape in the direction along the local magnetic field,

with an edge around  $\sim 3000$  km/s (Fig. 8d). The distribution function in regions (E) and (F) was further broadened in the direction along the magnetic field and had energy peaks in the low-energy range (2000–4000 km/s) rather than being a flat-topped shape (Figs. 8e and 8f). This broadening of the distribution function along the magnetic field may correspond to selective heating of electrons around the boundary between regions (D) and (E) mentioned above, for the temperatures of the cold proton component did not increase



**Fig. 8.** 1-D cuts of the electron velocity distribution function for the boundary crossing. Each panel (a–f) corresponds to each in Fig. 7. The PSDs in the perpendicular (blue) and parallel (green) directions to the local magnetic field are presented. Red curves show the one-count level of the detector, and the PSDs within a couple of red vertical lines may contain high contributions of the photoelectrons.

## Counter-streaming proton beams in the interior interaction region

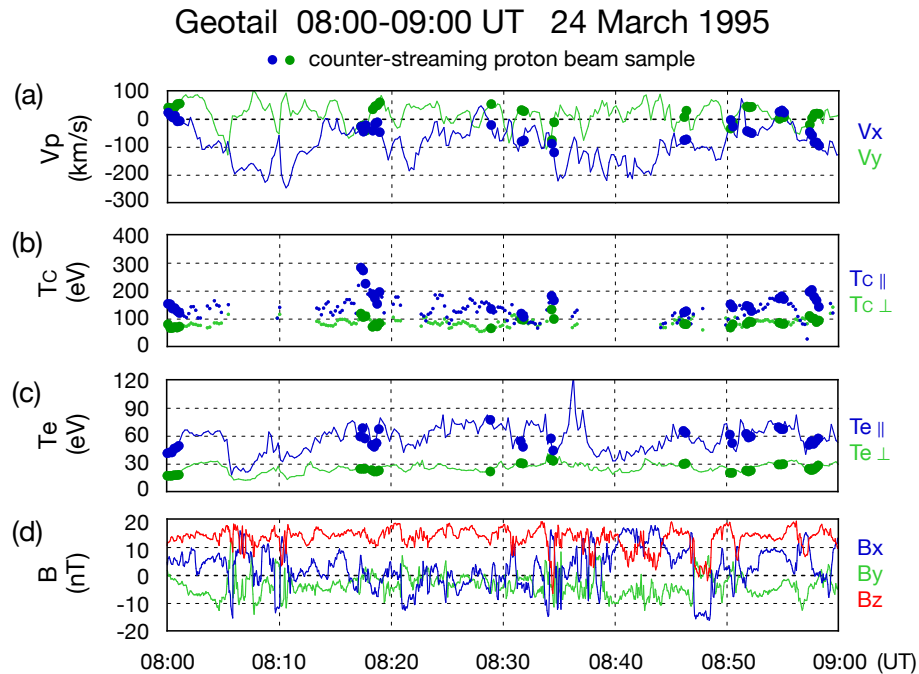


**Fig. 9.** Scatter plots in the interior interaction region between 07:16–09:10 UT. In Panels (a), (e), and (f), data for the counter-streaming protons are shown as black and red points, the latter of which show those obtained in region (D). In Panel (a) relation between  $N_p$  and  $V_x$  is presented, where gray points correspond to all other data obtained in the interior interaction region, and data in regions (A–F) are also plotted. The upper-right panel (b) shows the same data as shown in Panel (a) but color-coded by  $B_z$ . Panels (c) and (d) are color-coded by the parallel and perpendicular temperatures of electrons. In Panels (e) and (f), temperatures of the cold proton component and the electrons are presented, respectively, where gray points correspond to data obtained in the stagnant region ( $|V_x| < 100$  km/s).

as Geotail moved from region (D) into (E). In regions (E) and (F), the shape of the velocity distribution in the low-energy range resembled each other, but a high-energy component ( $> 15\,000$  km/s) was seen in region (F) while not in region (E). Fairfield et al. (2000) mentioned that electrons in the magnetosphere-like region of the vortical structures had a parallel anisotropy that is most often seen at 75 eV ( $\sim 5100$  km/s), which is near the energy of the low-energy bidirectional component in the present study.

#### 4.5 Survey through the interior interaction region

To further investigate counter-streaming proton beams found in region (D), we have surveyed the data through the interior interaction region between 07:16–09:10 UT. Apart from 6 data points obtained in region (D), 27 samples were identified as counter-streaming proton beams by visual inspection of proton PSDs.



**Fig. 10.** Counter-streaming proton beam samples observed between 08:00–09:00 UT. The large circles correspond to the counter-streaming proton beam samples. From the top, (a) proton velocity, (b) temperatures of the cold proton component (data in the stagnant region are shown), (c) electron temperatures, and (d) magnetic field are presented.

Figure 9a shows the relation between the proton density ( $N_p$ ) and the  $X$  component of the bulk proton flow ( $V_X$ ). As was mentioned by Hasegawa et al. (2006) and Takagi et al. (2006), the scatter plot of  $N_p$  and  $V_X$  is like a butterfly wing in shape. Most of the counter-streaming proton beam samples are seen in the stagnant region ( $|V_X| < 100$  km/s). The beam samples represent 15 percent of the data in the stagnant region.

In Fig. 9b the same data as shown in Fig. 9a are presented but color-coded by  $B_Z$ . Weak or negative  $B_Z$  (yellow-red colored points) are identified between the stagnant region and the magnetosheath-like dense region, which corresponds to the current layer in the vortical structures. The counter-streaming proton beam samples were not just at the current layer but on the magnetosphere side of the current layer, as well.

Figures 9c and d are the same as Figs. 9a and b but show the parallel and perpendicular temperatures of electrons, respectively, and in a color-coded fashion. The parallel electron temperature was enhanced in the stagnant region, while the perpendicular temperature was hardly enhanced there. The variation of the parallel temperature in the stagnant region was large, and the proton beam samples did not correspond to the highest parallel electron temperature, as shown later (Fig. 10).

Since most of the counter-streaming proton beam samples were obtained in the stagnant region, we now compare the beam samples and non-beam samples there. Figure 9e (9f)

presents the relation between the perpendicular and parallel temperatures of the cold proton component (the electrons) in the stagnant region ( $V_X < 100$  km/s). For the counter-streaming beam samples we use a two-Maxwellian fitting method; that is, the counter-streaming beams are dealt with as a simple cold proton component. The counter-streaming proton beam samples (black and red points) represent the majority to the right side of all data points, which means that the parallel anisotropy of the cold proton component is the strongest for the counter-streaming beam samples. In the plot of the electron temperatures, the counter-streaming proton beam samples appear not on the right side but in the midst of all data points. The strongest parallel anisotropy of the electrons corresponds to samples without the counter-streaming proton beams. This means that electrons receive further heating in the magnetosphere-like region, which is consistent with the enhancement in the electron temperatures in the regions (D–F) between 07:59:59–08:05:36 UT.

Figure 10 shows counter-streaming proton beam samples observed in the interval 08:00–09:00 UT. The large solid circles correspond to the counter-streaming proton beam samples. The parallel temperature of the cold proton component of the beam samples is enhanced (Fig. 10b), while the parallel temperature of electrons is not the highest (Fig. 10c). We note that the beam samples were found near the current layer that is characterized by the change in the magnetic field direction (Fig. 10d).



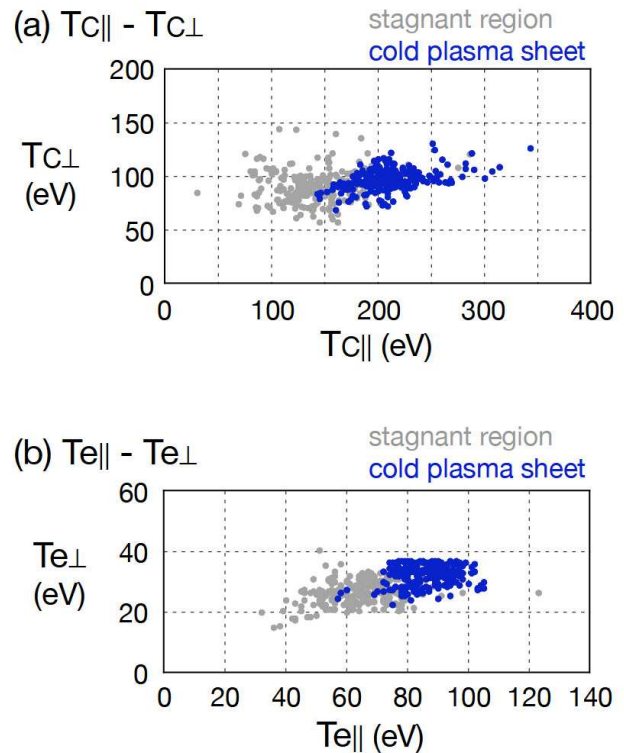
#### 4.6 Comparison with the cold plasma sheet

Let us compare the stagnant regions ( $|V_X| < 100$  km/s) inside the vortical structures with the cold plasma sheet between 09:11–09:59 UT. Figure 11 shows the perpendicular and parallel temperatures of the cold proton component (top panel) and electrons (bottom panel). The temperatures of the cold proton component in the cold plasma sheet ( $T_{C\perp} \sim 100$  eV and  $T_{C\parallel} \sim 210$  eV) were roughly 1.5 times higher than those in the stagnant regions, and in particular, the enhancement in the parallel temperature was remarkable. This means that the cold proton component would require further heating, predominantly along the magnetic field, in the cold plasma sheet or on the way from the vortices to the cold plasma sheet. A similar trend is seen in the electron temperatures. The averaged perpendicular and parallel temperatures of the electrons in the cold plasma sheet for the interval 09:11–09:59 UT were 32 eV and 86 eV, respectively, which are roughly 1.5 times higher than those in the magnetosphere-like regions in the KH vortical structure. The electron distribution function in the cold plasma sheet was flat-topped with an edge around 5000 km/s (Fig. 2d in Nishino et al., 2007b), being further broadened in shape along the magnetic field than that in regions (D–F). Enhancement of both electron and proton temperatures in the parallel direction suggests Fermi-like acceleration due to conservation of the second adiabatic invariant (discussed later).

## 5 Discussion

We first discuss a candidate mechanism for plasma heating inside the KH vortical structures (a summary is given in Fig. 12). In the present study, the sense of the proton anisotropy changed from perpendicular to parallel across the current layer. Around the border between regions (D) and (E), the low-energy portion of the proton distribution functions changed from counter-streaming beams to a single component with a strong parallel anisotropy. For the electrons, parallel heating was suggested in the following portions inside the vortical structure: (1) region (B) on the magnetosheath side of the current layer inside the KH vortical structure, (2) region (C) on the magnetosphere side of the current layer, and (3) in the magnetosphere-like regions (D–F). The heated electrons had a parallel anisotropy in the boundary layer and in the magnetosphere-like regions, which is consistent with previous studies (e.g. Traver et al., 1991; Fujimoto et al., 1998). Although bidirectional anisotropies of electrons were commonly seen through regions (B–F), the shape of the electron distribution functions in region (B) differs from that in regions (C–F). In particular, an enhancement of the bidirectional fluxes in the low-energy range (with speed less than 5000 km/s) was seen in regions (C–F) while not in region (B). Selective heating of electrons in the magnetosphere-like regions (D–F) is suggested by the

#### Stagnant region and cold plasma sheet

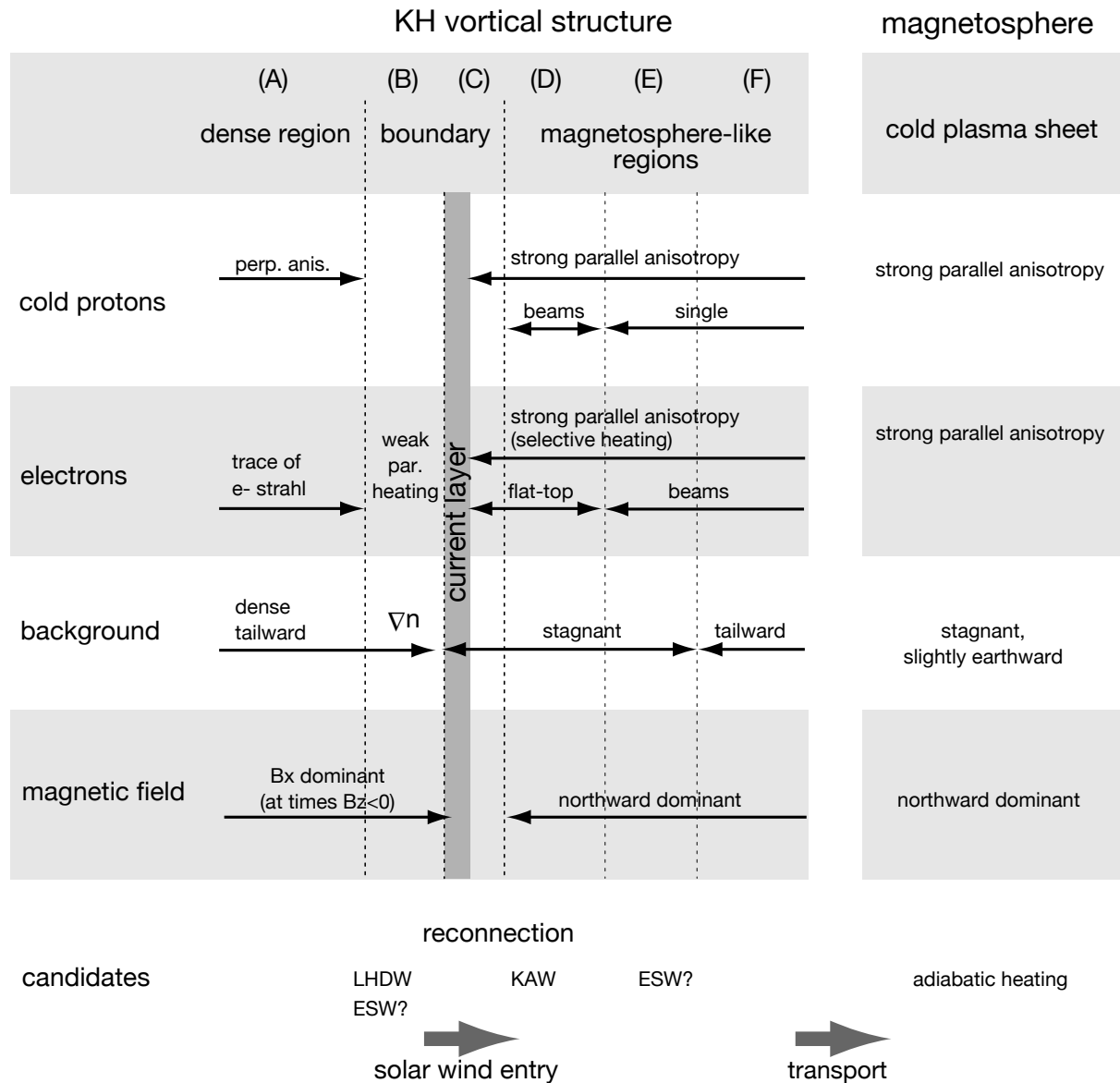


**Fig. 11.** Comparison between the temperatures in the stagnant regions of the interior interaction region and those in the cold plasma sheet.

fact that the electron temperatures increased while the proton temperatures did not there. Nishino et al. (2007b) reported that the parallel anisotropy of electrons is stronger than that of the cold proton component, which may be partly attributed to the selective heating of electrons inside the vortical structures.

The counter-streaming proton beams observed in region (D) may be evidence of magnetic reconnection inside the vortical structure. The proton beam speed was roughly comparable to the Alfvén speed, which is consistent with reconnection signatures. We examine whether the observed electron signatures fit the expectations of magnetic reconnection. The low-energy bidirectional electrons were observed in the magnetosphere-like regions inside the KH vortical structures. According to Hoshino et al. (2001), magnetic reconnection can accelerate the electron flow speed up to the electron Alfvén speed ( $V_{A,e}$ ), defined by  $V_{A,e} = B / \sqrt{\mu_0 m_e N_e} = V_A \sqrt{m_p / m_e}$ , where  $\mu_0$ ,  $N_e$ ,  $m_e$ , and  $m_p$  are the magnetic permeability, electron density, electron mass, and proton mass, respectively. When  $V_A$  was  $\sim 120$  km/s,  $V_{A,e}$  was  $\sim 5000$  km/s. The observed energy of the low-energy bidirectional electron beams is roughly comparable to  $V_{A,e}$ .





**Fig. 12.** Summary of observations and implicated candidates for physical mechanism.

Let us consider the reason why the observed proton PSDs do not represent a uni-directional beam but rather counter-streaming beams, for magnetic reconnection is likely to generate a uni-directional beam. One possibility is magnetic reconnection at two separate locations on the same field line (say, both above and below the spacecraft locations). Beams from two reconnection sites might lead to the counter-streaming velocity distributions. Another possibility is the mirror reflection of the beam due to a stronger magnetic field in the vortical structures. The mirror reflection would take place in the vortices, for the magnetic field strength in the vortices was at times enhanced to  $\sim 25$  nT while that of the counter-streaming beam samples was as small as  $\sim 15$  nT.

However, because the pitch angle of a charged particle that mirrors at the magnetic field strength 25 nT is as large as  $\sim 50^\circ$  for the location at the field strength 15 nT, mirror reflection inside the vortical structures might not play an important role in the generation of the counter-streaming beams.

We examine where in the vortex the counter-streaming proton beams existed, comparing observation and simulation studies. A candidate for the spacecraft trajectory relative to the vortex flow is illustrated in Fig. 13. The counter-streaming proton beam samples (in region (D)) seem to have existed near but not at the current layer, which implies that magnetic reconnection would take place in a region

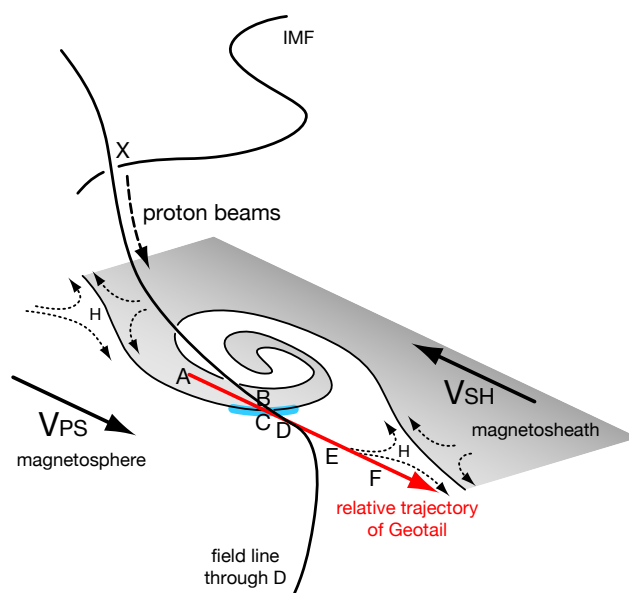
somewhat away from the satellite location. In a simulation study of the KH instability by Nakamura et al. (2006), magnetic reconnection takes place near the center between the two vortices. Since a hyperbolic point of the plasma flows (“H” in Fig. 13) exists when we are in the frame of reference moving with the vortices, a thinning of the current layer and the resultant magnetic reconnection may occur there. If the satellite was located on the magnetosphere side of the hyperbolic point, it might detect the counter-streaming protons that moved along the magnetic field line from the reconnection region(s). Furthermore, the observation of negative  $V_Y$  in the latter half of region (F) may represent plasma flow away from the hyperbolic point, and thus supports the idea that the spacecraft passed through the magnetosphere side of the hyperbolic point.

We discuss a possible heating mechanism around/inside the KH vortical structures, as well as in the cold plasma sheet. It has been proposed that magnetic reconnection in the KH vortices may excite kinetic Alfvén waves (KAWs) (Sibeck et al., 1999). The KAWs can heat electrons in the direction parallel to the local magnetic field and protons in the perpendicular direction (Lee et al., 1994; Johnson and Cheng, 1997; Johnson et al., 2001). The parallel anisotropy of electrons inside the KH vortical structure is consistent with signatures of the parallel heating by the KAWs. In addition, the ponderomotive force of the Alfvén waves can heat electrons but not sufficiently heat ions (Tsurutani et al., 2003). Since complicated interactions between the solar wind and the magnetospheric plasma inside/around the vortical structures may be a possible source of the Alfvén waves, the ponderomotive force may play an important role in selective heating of electrons in the magnetosphere-like region.

We suppose that the counter-streaming beams may excite specific waves that can heat electrons through the ion beam instability (Davidson et al., 1970). The observation that the parallel electron temperature of the proton beam samples was lower than that of non beam samples (Figs. 10b and c) suggests transformation of energy from the proton beams to the electrons via wave-particle interactions.

As a candidate for plasma transport processes across the magnetopause, diffusion mediated by the lower-hybrid drift wave (LHDW) has been proposed (Treumann et al., 1991; Shukla and Mamun, 2002). Obliquely propagating LHDWs can heat electrons in the direction parallel to the local magnetic field, giving rise to a flat-topped electron distribution function (Shinohara and Hoshino, 1999). While the study by Shinohara and Hoshino (1999) focuses on magnetic reconnection in the magnetotail, such processes might be of physical importance in the LLBL, as well. Furthermore, if magnetic reconnection occurs in the KH vortical structure, the accompanied LHDWs might play a role in parallel heating of the electrons.

Electrostatic solitary waves (ESWs) are expected to occur around the magnetopause and to form flat-topped electron distributions. Simulation studies showed that electron



**Fig. 13.** A schematic picture of the relative trajectory of the spacecraft in the frame of reference moving with the vortices. A candidate for the trajectory is presented as a red arrow on which rough locations of regions (A–F) are appended. The magnetic field line that passes through region (D) is depicted. The symbol “X” and “H” in the figure show candidates for reconnection region and the hyperbolic point, respectively. A blue-colored region between regions (B) and (C) corresponds to the observed current layer.

phase space holes created by the ESWs lead to particle trapping and heating (Omura et al., 1996; Goldman et al., 1999). Since recent observations by Cattell et al. (2002) showed that solitary waves exist at the magnetopause on the dayside under northward IMF, the observed flat-topped electrons distribution functions might be explained by the ESWs excited around the magnetopause. However, since the observations by Cattell et al. (2002) were performed not in the tail flank but on the dayside magnetopause, further inspections of the tail flank magnetopause are necessary in future works.

Comparing plasma inside the KH vortical structures with that in the cold plasma sheet, we have confirmed that the additional plasma heating must occur in the cold plasma sheet or on the way from the vortices to the cold plasma sheet. Nishino et al. (2007a,b) mentioned that some portion of the observed parallel anisotropies of the electrons and the cold proton component may be attributed to adiabatic heating due to earthward convection in the plasma sheet under conservation of the first and second adiabatic invariants. The idea of adiabatic heating can also be supported by the fact that increase ratios of temperatures of the electron and the cold proton component from the KH vortices to the cold plasma sheet resembled each other (roughly  $\sim 1.5$ ). We also suggest that the cold plasma observed in the plasma sheet would come into the magnetosphere across the downtail magnetopause,

possibly via highly developed KH vortices. Since the KH vortical structure is expected to develop in the tail regions more than near the terminator (Hasegawa et al., 2006), efficiency of plasma transport by magnetic reconnection inside the KH vortices should be higher as the distance downtail increases. This qualitative picture is consistent with the spatial distribution of the cold plasma obtained from low altitude satellite data mapped into the tail, as reported by Wing et al. (2005).

## 6 Conclusions

We have studied temperature anisotropies for a Kelvin-Helmholtz (KH) vortex event on the duskside. The cold proton component had a strong parallel anisotropy in the magnetosphere-like region, and it consisted of counter-streaming beams in the region near the boundary, which suggests magnetic reconnection inside the KH vortical structure. The parallel anisotropy of electrons starts to be enhanced on the magnetosheath side of the current layer in the KH vortical structure, and low-energy bidirectional electron beams or flat-topped electron distribution functions in the direction along the local magnetic field are apparent on the magnetosphere side of the current layer. We have suggested the occurrence of selective heating of electrons inside the vortical structure via wave-particle interactions. Further heating predominantly along the magnetic field by the Fermi-like acceleration is likely to be taking place in the cold plasma sheet or on the way from the vortices to the cold plasma sheet.

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